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# The Herschel view of GAS in Protoplanetary Systems (GASPS). First comparisons with a large grid of models

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LETTER TO THE EDITOR

# The *Herschel* view of Gas in Protoplanetary Systems (GASPS)<sup>\*</sup>

## First comparisons with a large grid of models

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### ABSTRACT

The *Herschel* GASPS Key Program is a survey of the gas phase of protoplanetary discs, targeting 240 objects which cover a large range of ages, spectral types, and disc properties. To interpret this large quantity of data and initiate self-consistent analyses of the gas and dust properties of protoplanetary discs, we have combined the capabilities of the radiative transfer code MCFOST with the gas thermal balance and chemistry code ProDiMo to compute a grid of  $\approx 300\,000$  disc models (DENT). We present a comparison of the first *Herschel*/GASPS line and continuum data with the predictions from the DENT grid of models. Our objective is to test some of the main trends already identified in the DENT grid, as well as to define better empirical diagnostics to estimate the total gas mass of protoplanetary discs. Photospheric UV radiation appears to be the dominant gas-heating mechanism for Herbig stars, whereas UV excess and/or X-rays emission dominates for T Tauri stars. The DENT grid reveals the complexity in the analysis of far-IR lines and the difficulty to invert these observations into physical quantities. The combination of *Herschel* line observations with continuum data and/or with rotational lines in the (sub-)millimetre regime, in particular CO lines, is required for a detailed characterisation of the physical and chemical properties of circumstellar discs.

**Key words.** Astrochemistry; circumstellar matter; protoplanetary discs ; stars: formation; Radiative transfer; Methods: numerical; line: formation

## 1. Introduction

The dust phase of circumstellar discs has received a lot of attention in the last few decades, giving us a clearer picture of their structure and dust content through many continuum surveys in various wavelengths regimes (e.g. Beckwith et al. 1990; Andrews & Williams 2007; Evans et al. 2007), complemented by detailed studies of individual objects, combining spectral energy distributions (SEDs) and resolved maps in scattered light and thermal emission (e.g. Pinte et al. 2008; Duchêne et al. 2010).

Although gas represents 99% of the initial mass of discs, it has been more difficult to observe and is mostly restricted to millimetre lines probing the cold outer disc, where the freeze-out of molecules is important (e.g. Dent et al. 2005; Schaefer et al. 2009), and near-IR lines which are only emitted from the hot inner parts of discs (e.g. Najita et al. 2003; Brittain et al. 2007). The high sensitivity of *Herschel* (Pilbratt et al. 2010) opens an opportunity to systematically probe the gas phase of discs, in particular the warm atomic and molecular layer responsible for the bright gas emission lines in the far-IR. The GASPS open time key program (see Dent et al., in prep. and Mathews et al. 2010) is a large survey of gas in discs with a gas mass sensitivity com-

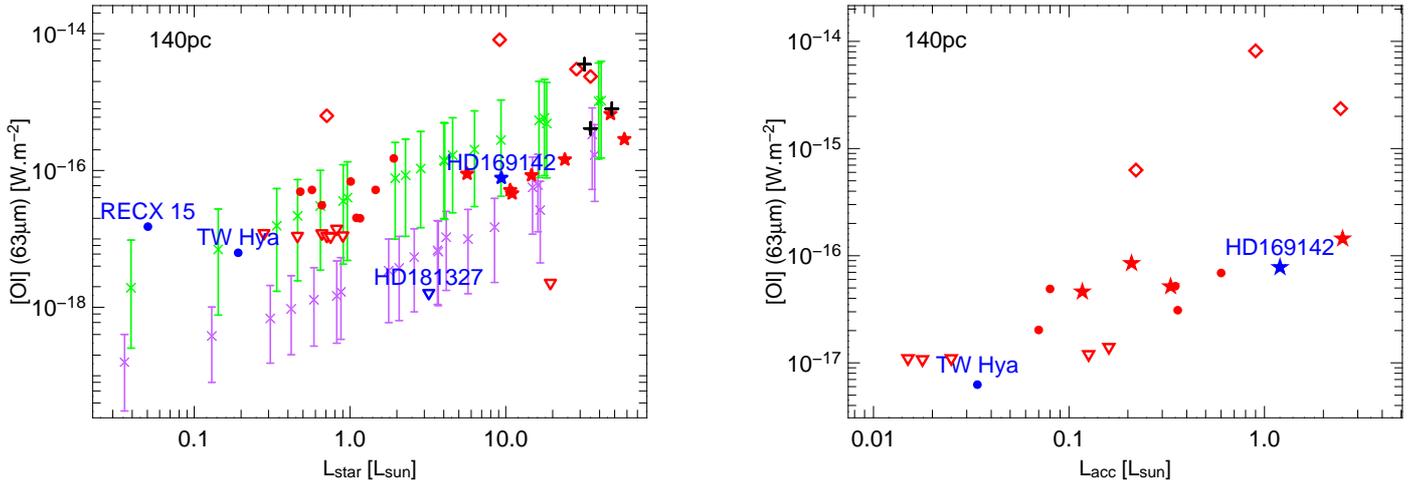
parable to the dust surveys. GASPS will observe several atomic and molecular lines in about 240 protoplanetary disc systems with ages in the critical 1 to 30 million year age range during which planets form and the gas seems to dissipate.

The interpretation of gas observations is complicated by the large number of processes at play: processing of radiation by dust grains, disc thermal structure, chemistry, excitation and destruction of molecules, freeze-out and desorption on the dust grains, etc. To estimate the relative importance of these mechanisms as a function of age, stellar properties, disc structure, and dust content, we have computed a large grid of synthetic SEDs and gas emission lines, named Disc Evolution with Neat Theory (DENT, Woitke et al. 2010). Here, we confront the trends identified in the DENT grid with the first GASPS observations.

## 2. The DENT grid and initial GASPS data

The DENT grid is intended as a statistical tool to investigate the influence of stellar, disc, and dust properties on the various continuum and line observables, and to study to what extent these dependencies can be inverted to retrieve disc properties. The grid relies on the combined capabilities of the 3D radiative transfer code MCFOST (Pinte et al. 2006, 2009) and the gas thermal balance and chemistry code ProDiMo (Woitke et al. 2009; Kamp et al. 2010). Spectral energy distributions and line fluxes of [OI], [CII], <sup>12</sup>CO, ortho-H<sub>2</sub>O and para-H<sub>2</sub>O are predicted for

<sup>\*</sup> *Herschel* is an ESA space observatory with science instruments provided by European-led Principal Investigator consortia and with important participation from NASA.



**Fig. 1.** [OI]  $63\mu\text{m}$  line flux as a function of the stellar luminosity (left panel) and accretion luminosity (right panel). Red and blue points, stars, triangles, and diamonds represents GASPS observations. Points are T Tauri stars while stars are Herbig stars. Diamonds are sources with extended emission, which could originate from an outflow or envelope. Triangles are GASPS upper limits. Names are only indicated for SDP sources (blue). Black pluses are ISO observations. Observations were scaled to a distance of 140 pc. The crosses show the median values of the DENT models. The full vertical lines represent the 1-sigma range of models. Only models with a gas mass larger than  $10^{-4} M_{\odot}$  are plotted. Green and purple points are high- and low-UV models respectively ( $f_{\text{UV}} = 0.1$  and  $0.001$ ). Accretion luminosities from Hartmann et al. (1998) and Garcia Lopez et al. (2006).

more than 300 000 discs models. The DENT grid was built by systematically exploring an 11-dimension parameter space (see Woitke et al. 2010, table 1). In particular the DENT grid explores the effect of varying the central star (age, mass, UV excess), disc dust mass and gas-to-dust mass ratio, inner and outer radii, flaring and surface density exponents, grain sizes, and presence of dust settling. It is important to keep in mind that even with the large number of calculated models, the sampling of each parameter remains coarse and that the DENT grid does not reflect the statistics of objects in GASPS (each parameter has been sampled uniformly and not following the distributions of the GASPS target list). We refer the reader to Woitke et al. (2010) for details about the grid properties and computational implementation.

We include here data obtained during the science demonstration phase (SDP, Mathews et al 2010), as well as GASPS data reduced prior to 2010 April 23, which will be presented in detail in following papers. Due to the limited number of sources, statistical analyses remain premature, but initial comparisons with the model predictions are necessary to ensure that the range of models cover the GASPS observations.

### 3. Results and discussion

One of the main reasons to compute the DENT grid was to estimate the degeneracies between parameters, i.e. how they influence the various lines and how far *Herschel* line observations can be inverted to assess the physical and chemical conditions of the disc. Not surprisingly, the DENT grid revealed that many parameters affect the predicted line fluxes and SEDs, and degeneracies between parameters are common and complex, which makes the interpretation of lines fluxes difficult.

#### 3.1. Gas heating processes

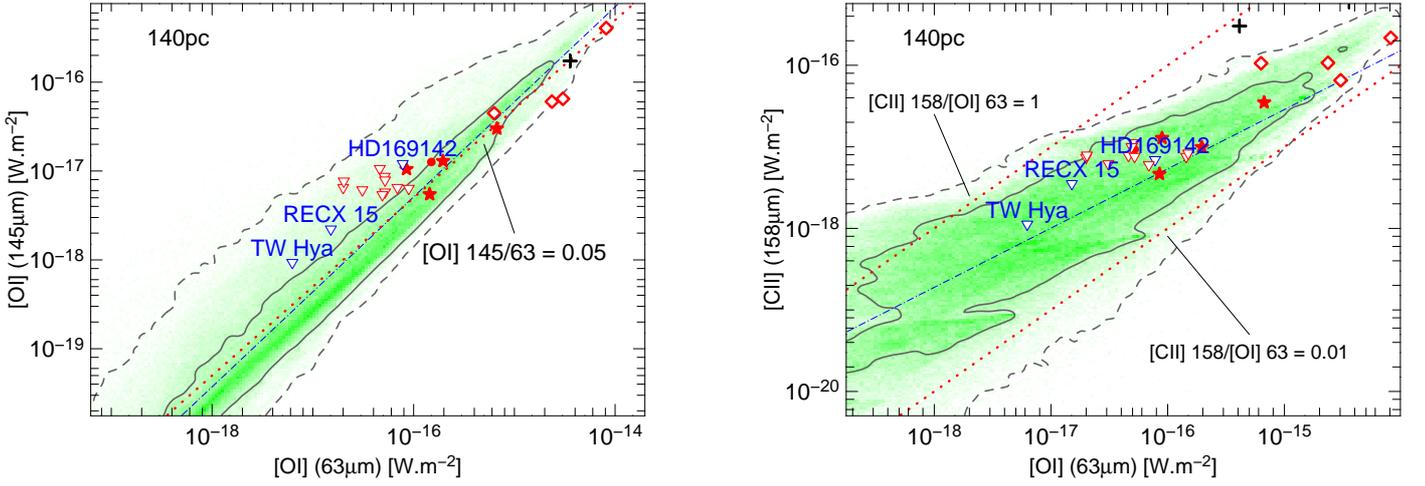
Figure 1 plots the [OI]  $63\mu\text{m}$  line flux as a function of the stellar luminosity and accretion luminosity. The DENT grid predicts a correlation between the line flux and the stellar luminosity. All

observational points (except sources with a large outflow or envelope) lie within the 1-sigma envelope of the models.

For T Tauri stars, the detected line fluxes are well reproduced by models with a high UV excess ( $f_{\text{UV}} = 0.1$ , see Woitke et al. 2010), suggesting that UV emission produced by accretion onto the star is one of the main gas-heating processes. The right panel of Fig. 1 indeed suggests a trend between [OI] line flux and accretion luminosity. On the other hand, Herbig Ae/Be stars show a strong correlation between the line flux and stellar luminosities, with a much smaller scatter than for T Tauri stars. Large UV excesses do not seem necessary to reproduce the Herbig observations (data points lie between models with high- and low-UV excess). This suggests that stellar radiation is the dominant gas-heating source for Herbig stars. Because these sources radiate large photospheric UV emission, the accretion luminosity represents a smaller fraction of the UV luminosity and is not as critical a gas heating mechanism as for T Tauri stars. This is also consistent with the small fraction of large accretors ( $L_{\text{acc}} > 0.1 L_{\odot}$ ) among Herbig stars (Garcia Lopez et al. 2006).

X-ray irradiation, which is not yet included in the DENT grid, can also contribute significantly to the gas heating and chemistry for low-mass objects (e.g. Glassgold et al. 2004; Semenov et al. 2004; Meijerink et al. 2008; Hollenbach & Gorti 2009; Ercolano & Owen 2010) and higher fluxes can be expected for sources with typical T Tauri X-ray emission. The small number of sources observed by GASPS so far prevents us from distinguishing between UV and X-rays for the main heating process for low mass objects. The full GASPS survey should provide detailed answers on these aspects.

Figure 2 presents the [OI]  $145\mu\text{m}$  and [CII]  $158\mu\text{m}$  line fluxes as a function of the [OI]  $63\mu\text{m}$  line. The models are in excellent agreement with the GASPS observations. The DENT grid predicts a correlation between the [OI] 63 and  $145\mu\text{m}$  line fluxes. A regression fit of all the DENT points indicates that both line fluxes are almost proportional ( $f([\text{OI}] 145\mu\text{m}) \propto f([\text{OI}] 63\mu\text{m})^{0.98}$ ), with a [OI] 145/63 line ratio around 0.05 on average. The presence of scatter in the plot illustrates the wide range of physical conditions encountered in the DENT



**Fig. 2.** Correlations between line fluxes. Symbols as in Fig. 1. *Left panel:* [OI] 145  $\mu\text{m}$  as a function of [OI] 63  $\mu\text{m}$ . *Right panel:* [CII] 158  $\mu\text{m}$  as a function of [OI] 63  $\mu\text{m}$ . Small green points: DENT models. The full and dashed grey contours represent the regions that contain 68 % and 99.7 % of the models respectively. The blue dot-dashed lines represent power law regression fits to the entire DENT grid. The red dotted lines correspond to constant line ratios. In the left panel the line corresponds to the average ratio in the DENT grid: 0.05. This ratio corresponds to the limit case for optically thick lines at temperatures higher than a few hundred Kelvin (Tielens & Hollenbach 1985).

grid. The correlation, however, is very strong (Pearson correlation coefficient of 0.97) and holds for several orders of magnitude in line flux. These results agree with the prediction of Tielens & Hollenbach (1985) for a 1D photodissociation region (see their Fig. 2). Our detailed line modelling of the 2D PDR disc surface with varying density and irradiation confirms the picture drawn from these 1D models: the oxygen lines are optically thick and originate in a relatively high-temperature gas ( $\geq 100$  K). As a consequence, this line ratio does not provide constraints on the local gas density and temperature in most cases.

Deviations from this average ratio of 0.05 are interesting though. In particular, two Herbig Ae with known outflows observed by GASPS present a small line ratio around 0.025. According to Tielens & Hollenbach (1985), this value cannot be obtained for optically thick lines. This suggests that a significant fraction of the line fluxes originates from an optically thin region above the bulk of the disc (potentially the outflow) with a temperature between 40 and 200 K. These sources will be studied in detail in following GASPS papers.

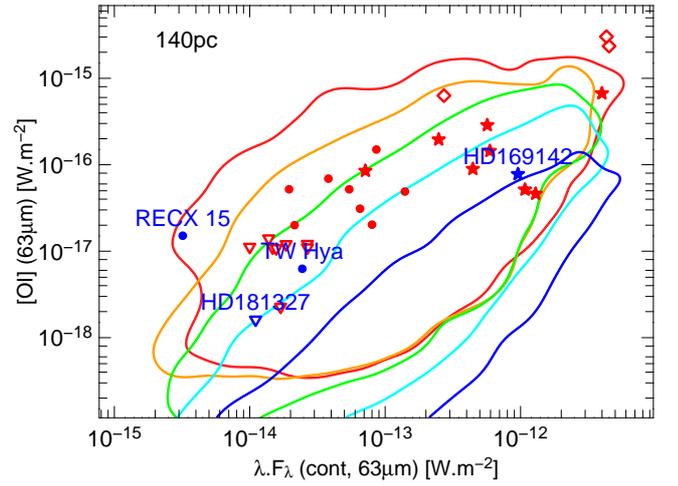
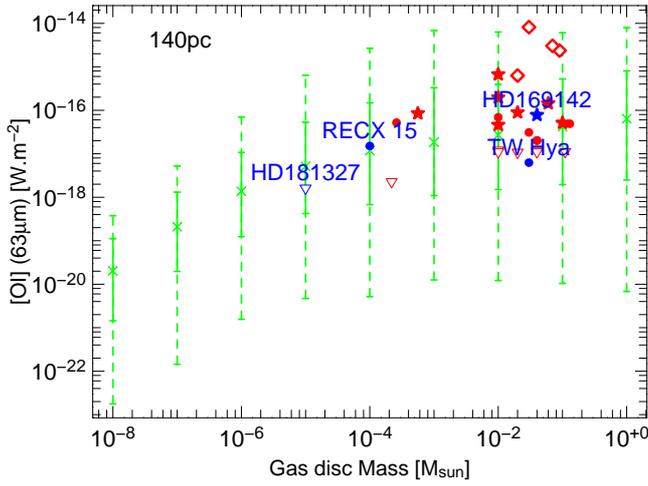
The [CII] 158  $\mu\text{m}$  line also presents a correlation with the [OI] 63  $\mu\text{m}$ , but with a much larger scatter. A regression fit of the DENT models indicates that on average  $f(\text{[CII] } 158 \mu\text{m}) \propto f(\text{[OI] } 63 \mu\text{m})^{0.7}$ , with a decreasing line ratio as the line flux increases. This suggests an increasing gas temperature (Fig. 3 in Tielens & Hollenbach 1985) in the disc towards more luminous objects. Most of the DENT models lie in a region where the line ratio is between 0.01 and 1, suggesting gas temperatures higher than 100 K. More detailed analyses of the [CII] line are complicated by several factors: our disc models show that the line originates from larger radii and lower density regions than the [OI] lines, and it is very sensitive to the amount of UV radiation.

### 3.2. Gas mass and gas-to-dust mass ratio

In addition to the main trend of an increasing line flux with (UV) luminosity, there is also a trend with gas mass, where the synthetic line flux increases with mass, but the correlation seems to saturate above  $10^{-4} M_{\odot}$  (see also Woitke et al. 2010), preventing direct inversion of the line flux into a gas mass (Fig. 3).

The right panel of Fig. 3 shows the correlation between the [OI] 63  $\mu\text{m}$  line flux and the adjacent continuum. Because the line and continuum emissions are optically thick in most cases, these fluxes give an indication of the relative temperatures and projected surface area of the emitting regions (gas and dust). As the stellar luminosity increases, the region of the disc which is warm enough to contribute significantly to the emission also increases, resulting in larger fluxes. This behaviour is observed for most GASPS sources where the line flux roughly increases with the continuum level, but with a significant scatter. As a consequence, this indicates that a large fraction of discs with a significant far-IR excess will be detected in [OI] by *Herschel*. The separation of the DENT models according to their gas-to-dust mass ratio suggests that most objects are gas-rich (gas/dust mass ratio  $> 10$ ). The large scatter in the models, and the optical depth in the continuum and the line, precludes however, in most cases, a precise estimate of the gas-to-dust ratio for individual sources. For instance, no direct ratio (or upper limit for HD 181327) can be estimated from this diagnostic alone for HD 169142 and TW Hydra, for which the line fluxes can be reproduced by any ratio between 1 and 1000. In addition, the contribution of an outflow to the line flux may affect the estimation of the disc gas-to-dust ratio and needs to be accounted for.

Greater observational constraints and more detailed modelling is required to estimate the gas mass and gas-to-dust ratio. In particular, the combination of low rotational level transitions of CO with oxygen lines offers a valuable proxy to estimate the amount of gas in discs. Figure 4 plots the  $^{12}\text{CO } J=3 \rightarrow 2$  line flux as a function of the [OI] 63  $\mu\text{m}$  line flux. We stress that the accuracy in the calculated CO abundances is limited by our approximate treatment of self-shielding (see Woitke et al. 2009), but this does not affect our conclusions. For low-mass discs, this diagram allows a clear distinction of the gas disc mass. As the mass increases, lines become optically thick and the corresponding fluxes saturate, preventing determination of the gas mass. Current CO surveys can only reach sources in this saturation regime (see for instance data from Dent et al. 2005 in Fig. 4), but this perspective is particularly interesting in the context of *ALMA*, which will offer high sensitivity for CO lines



**Fig. 3.** [OI] 63  $\mu\text{m}$  line flux as a function of the gas mass (left) and adjacent continuum (right). Symbols as in Fig. 1. The full and dashed vertical lines represent the 1- and 3-sigma range of models. The values for the gas mass of the observed sources are *indicative* only, they were estimated from millimetre emission and assuming a gas-to-dust ratio of 100. On the right panel, contours represent the regions that contain 68 % of the models as a function of the gas-to-dust mass ratio: red = 1000, orange = 100, green = 10, light blue = 1, dark blue = 0.1.

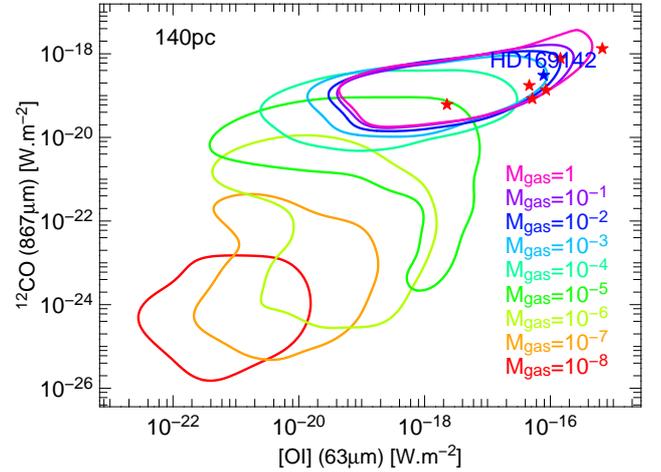
( $\approx 10^{-23} \text{ W.m}^{-2}$ ). Similar diagrams combining  $^{13}\text{CO}$  (not included in DENT, but see Meeus et al 2010 and Thi et al 2010),  $\text{C}^{18}\text{O}$  and [OI] 145  $\mu\text{m}$ , which saturate at higher masses due to lower optical depths, will further help to overcome this degeneracy. As oxygen lines are sensitive to warm gas in the inner 10-30 AU (for T Tauri stars), they offer complementary views to the low-J CO lines which probe regions outside of 20-40 AU, especially when resolved maps of the CO emission are available.

#### 4. Summary and conclusions

The GASPS survey will offer unique views of the gas and dust phases of protoplanetary discs. In order to provide statistical tools to help the interpretation of the survey results, we interfaced the MCFOST and ProDiMo codes and calculated a large grid of models sampling the range of discs observed by GASPS. This will allow us to determine some of the physical conditions within discs. The initial results from the GASPS survey tend to confirm the predictions of the DENT grid, illustrating the main parameters affecting the line fluxes, namely the UV excess and/or X-ray emission for T Tauri stars and the UV stellar irradiation for Herbig stars. This is a highly relevant point to be considered in subsequent open time programs on discs.

The interpretation of line results remains difficult and their inversion into physical parameters must be performed with caution, because the DENT grid highlights considerable degeneracies between parameters and the complex interplay between various physical processes. The [OI] 63  $\mu\text{m}$  is crucial for breaking some of the degeneracies. By combining this line with continuum and/or (sub)mm rotational lines, we can possibly distinguish various parameters. Meeus et al (2010) and Thi et al (2010) illustrate how far this inversion can be performed when high quality data sets with a wide range of observational techniques are available.

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**Fig. 4.**  $^{12}\text{CO}$  J=3 $\rightarrow$ 2 (867  $\mu\text{m}$ ) as a function of [OI] 63  $\mu\text{m}$ . Contours represent the regions that contain 68 % of the models as a function of the gas mass (in solar masses). Data points from Dent et al. (2005).

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